



PHANGS–JWST First Results: Massive Young Star Clusters and New Insights from JWST Observations of NGC 1365

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Abstract

A primary new capability of JWST is the ability to penetrate the dust in star-forming galaxies to identify and study the properties of young star clusters that remain embedded in dust and gas. In this Letter we combine new infrared images taken with JWST with our optical Hubble Space Telescope (HST) images of the starbursting barred (Seyfert2) spiral galaxy NGC 1365. We find that this galaxy has the richest population of massive young clusters of any known galaxy within 30 Mpc, with ~ 30 star clusters that are more massive than $10^6 M_{\odot}$ and younger than 10 Myr. Sixteen of these clusters are newly discovered from our JWST observations. An examination of the optical images reveals that 4 of 30 ($\sim 13\%$) are so deeply embedded that they cannot be seen in the Hubble *I* band ($A_V \gtrsim 10$ mag), and that 11 of 30 ($\sim 37\%$) are missing in the HST *B* band, so age and mass estimates from optical measurements alone are challenging. These numbers suggest that massive clusters in NGC 1365 remain completely obscured in the visible for $\sim 1.3 \pm 0.7$ Myr and are either completely or partially obscured for $\sim 3.7 \pm 1.1$ Myr. We also use the JWST observations to gain new insights into the triggering of star cluster formation by the collision of gas and dust streamers with gas and dust in the bar. The JWST images reveal previously unknown structures (e.g., bridges and overshoot regions from stars that form in the bar) that help us better understand the orbital dynamics of barred galaxies and associated star-forming rings. Finally, we note that the excellent spatial resolution of the NIRCAM F200W filter provides a better way to separate barely resolved compact clusters from individual stars based on their sizes.

Unified Astronomy Thesaurus concepts: Barred spiral galaxies (136); Young star clusters (1833); Young massive clusters (2049)

1. Introduction and Motivation

The discovery that massive young super star clusters (SSCs) are forming in merging and starbursting galaxies was one of the earliest results from the Hubble Space Telescope (e.g., Holtzman et al. 1992; Whitmore et al. 1993; O’Connell et al. 1994;



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Barth et al. 1995). This discovery provided a direct observational window into the conditions needed to form massive clusters, some of which will likely evolve into globular clusters (e.g., Ashman & Zepf 1992; Elmegreen & Efremov 1997; Kruijssen 2015). Many of these SSCs form in dusty regions, including luminous infrared galaxies (e.g., Adamo et al. 2020a; Linden et al. 2021) with radio observations identifying some of the youngest, most embedded ones (e.g., Sandqvist et al. 1995; Whitmore & Zhang 2002; Johnson & Kobulnicky 2003; He et al. 2022).

The dusty, barred spiral galaxy NGC 1365, at a distance of 19.57 ± 0.78 Mpc (Anand et al. 2021), has one of the highest rates of star formation ($16.9 M_{\odot} \text{ yr}^{-1}$) of any galaxy within 20 Mpc (the next highest in the PHANGS–HST sample of 38 galaxies is $7.6 M_{\odot} \text{ yr}^{-1}$; Lee et al. 2022) and is known to have recently (past ≈ 10 Myr) formed very massive ($\gtrsim 10^6 M_{\odot}$) clusters (e.g., Kristen et al. 1997; Galliano et al. 2008).

The main goal of this Letter is to use new infrared imaging of NGC 1365, taken with the recently launched JWST, plus our optical Hubble Space Telescope (HST) images to collect a complete census of these very young massive clusters, with a focus on discovering clusters missed by previous optical studies because they are enshrouded in dust.

The rest of this Letter is organized as follows. Section 2 describes the observations and photometric reductions. Section 3 examines various objects in Region 1, the northern half of the central star-forming ring, and compiles a census of massive star clusters. Section 4 examines some interesting features in Region 2, just outside the central region along the eastern bar. Our main findings are summarized in Section 5.

2. Observations and Reductions

The primary data sets used in this paper are from the PHANGS–JWST Treasury Proposal (PI: J.C. Lee, GO-2107) and the PHANGS–HST project (PI: J.C. Lee, GO-15654).

PHANGS–HST is a Cycle 26 HST Treasury program that obtained imaging for 38 nearby spiral galaxies in the following five filters: F275W (NUV), F336W (*U*), F438W (*B*), F555W (*V*), and F814W (*I*), with the Wide Field Camera 3 (or Advanced Camera for Surveys in cases where archival data were available). See Lee et al. (2022) for details about PHANGS–HST.²⁷

PHANGS–JWST is a Cycle 1 JWST Treasury program that obtained imaging for the 19 (out of 38) PHANGS–HST galaxies that have integral field unit (IFU) maps obtained with the Very Large Telescope (VLT)/MUSE (Emsellem et al. 2022). Atacama Large Millimeter/submillimeter Array (ALMA) CO(2–1) observations are available for a larger set of 90 PHANGS–ALMA (Leroy et al. 2021) galaxies that include all 19 of the PHANGS–JWST subset. The JWST observations were made in the NIRCAM filters F200W, F300M, F335M, and F360M and in the MIRI filters F770W, F1000W, F1130W, and F2100W. See Lee et al. (2023) for additional details about the observations and basic reductions.

Aperture photometry was performed using a $0''.155$ (i.e., 5 pixel) radius aperture with a sky annulus between $0''.217$ and $0''.310$, as described in Rodriguez et al. (2023). This subtends a five pixel radius for the F200W NIRCAM observations and includes approximately 75% of the light from a point source, but smaller percentages at longer wavelengths (e.g., $\sim 40\%$ for

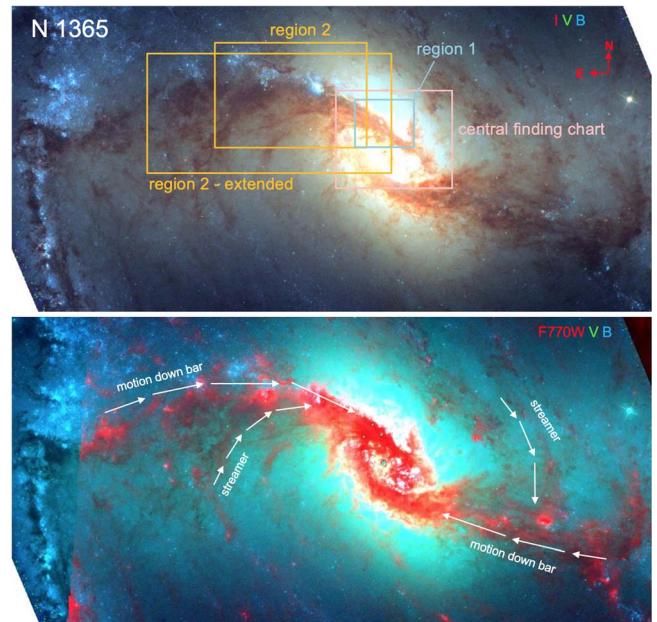


Figure 1. An optical *IVB* color image taken by Hubble of NGC 1365 is shown in the upper panel, and a JWST infrared (F770W) plus Hubble (*V* and *B*) image is shown in the bottom panel. The full field of view covers an area of roughly $80'' \times 160''$ ($7.5 \text{ kpc} \times 15.0 \text{ kpc}$). The fields of view for regions discussed in the text are shown, as well as the motion along the bar and along two of the streamers. We note that the extensive star formation in the spiral arm and the end of the eastern bar on the left edge of the image is outside the F770W field of view.

the F1000W filter and $\sim 10\%$ for F2100W). For this reason, the JWST magnitudes are used in a relative rather than absolute sense throughout this Letter. Magnitudes from Hubble are referred to as NUV, *U*, *B*, *V*, and *I* and those from JWST are referred to by their filter names in order to avoid confusion. Source selection was performed based on the F335M image, and the corresponding positions were then measured in the other bands. The ABmag system is used throughout.

In the future, aperture corrections will be determined for both stars and clusters. A variety of other observational parameters will also improve with time, including image alignment and geometric solutions, photometric zero-points, and background subtraction, as discussed in Lee et al. (2023). For these reasons the data and massive cluster catalog we compile here should be considered preliminary. A more definitive treatment will be provided in future papers.

3. Region 1: Massive Young Star Clusters

Figure 1 shows the locations of several parts of NGC 1365 that will be discussed in this Letter. Our primary focus will be on the central $28'' \times 26''$, or $2.7 \text{ kpc} \times 2.5 \text{ kpc}$, at the assumed distance of 19.57 Mpc, denoted as the “central finding chart” in Figure 1 and shown in more detail in Figure 2. Region 1 is in the northern part of this region, where some of the most massive young clusters in the nearby universe have formed (Galliano et al. 2008, 2012).

3.1. A Compilation of Massive Young Clusters in NGC 1365

In order to build a complete census of massive, young clusters in NGC 1365, we start by identifying all optically selected clusters that have estimated ages younger than 10 Myr and masses higher than $10^6 M_{\odot}$, from the cluster catalog and

²⁷ <https://archive.stsci.edu/hlsp/phangs-hst>

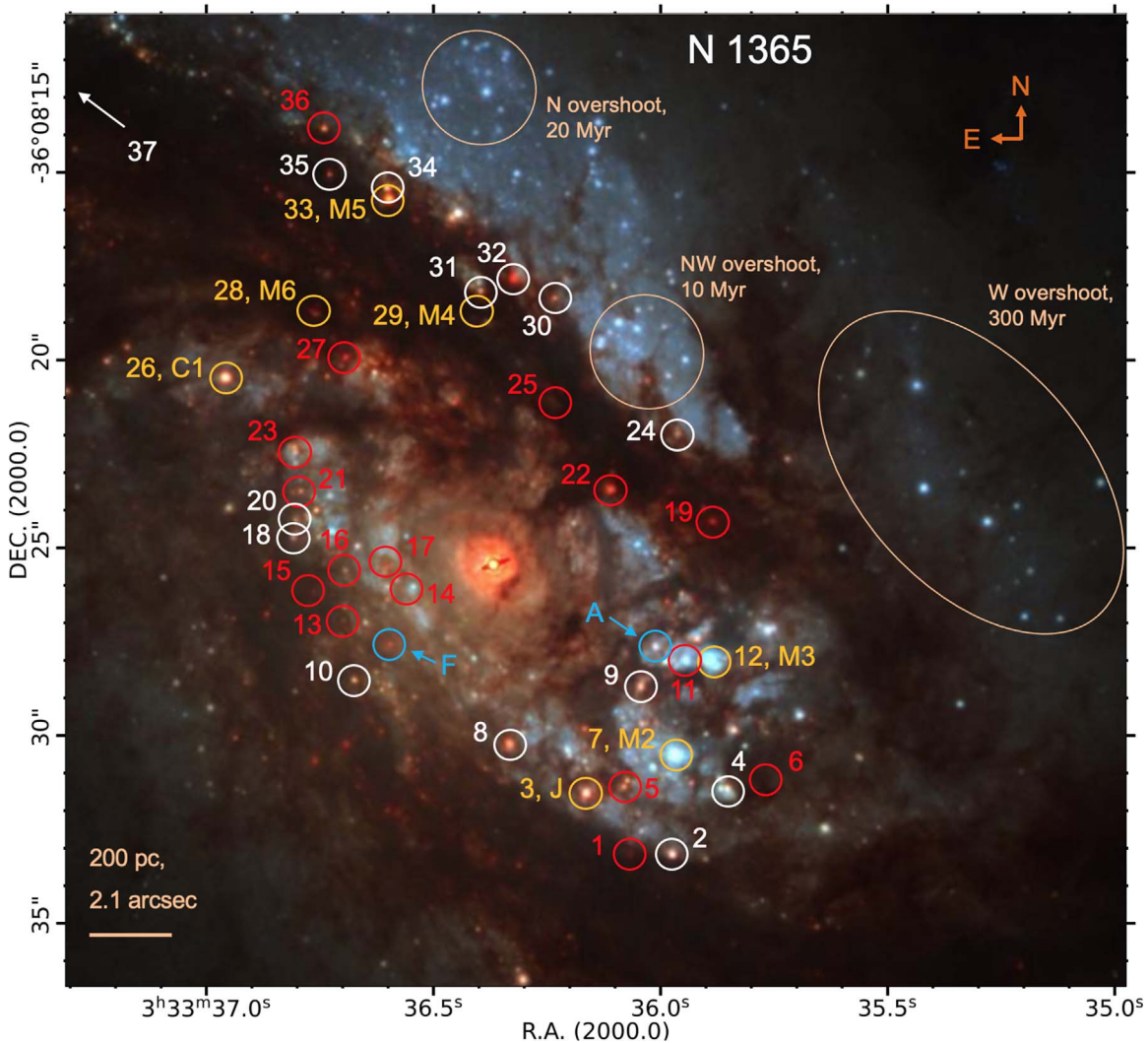


Figure 2. *I/B* Hubble image of the central region of NGC 1365 showing the location of the massive young star clusters from Table 1. Note that 36 of the 37 objects are in this central region, and only one (37) is slightly farther up the dust lane along the bar. White circles show the clusters from the Hubble compact clusters catalog. “Historical” clusters are identified by the gold circles. Red circles show the new clusters discovered in the JWST images. The locations of radio continuum sources “F” (believed to be related to the AGN jet) and “A” (believed to be related to a supernova remnant; see Sakamoto et al. 2007) are circled in blue, but are not included in the 37 clusters listed in Table 1. Three overshoot regions discussed in Section 4.2 are also identified.

age-dating analysis based on multiband HST optical images presented in Turner et al. (2021) and Whitmore et al. (2023). Three clusters are removed from the original list since they have very faint F2100W magnitudes indicative of old populations, unlike the rest of the young clusters in our final sample that are all very bright in F2100W. We find 19 clusters that satisfy our criteria (these are the 19 clusters in Table 1 that have Cluster IDs in column (2)).

To this list, we add two “historical” clusters (M4 and M6) that were previously discovered in radio continuum images (e.g., Sandqvist et al. 1982, 1995), with estimated ages less than 10 Myr and masses in the range 3×10^6 to $10^7 M_{\odot}$ (Galliano et al. 2008, 2012). These are comparable to the most massive, recently formed clusters in the merging Antennae (Whitmore et al. 2010) and NGC 3256 (Zepf et al. 1999; Mulia et al. 2016; Adamo et al. 2020b) galaxies.

Finally, we add 16 new, partly or heavily embedded massive young clusters identified in the new JWST images. These have been selected to be brighter than 20.0 mag in at least one of the three filters that contain polycyclic aromatic hydrocarbon (PAH) emission (F335M, F770W and F1130W). As shown in

Table 1, many sources meet this criterion in all three filters, but a few only qualify in one or two.

We determined the magnitude limit of 20.0 mag to be included in our sample by requiring that the 19 clusters detected from HST (with one or two exceptions) be included. As will be discussed in Section 3.2.1, we also require the F200W magnitude to be brighter than 20.0 mag to ensure there is a bright stellar continuum source present, rather than just a dense cloud of gas/dust energized by nearby sources. We also require that the concentration index (CI; defined as the magnitude difference measured in 1 and 4 pixel aperture radii) in the F200W band to be larger than that of a point source: $CI_{F200W} > 1.4$ mag.

Different cluster selection criteria are possible, and in fact several other works within the PHANGS consortium have selected clusters using a single filter: the F335M filter (Rodríguez et al. 2023), the F1000W filter (Schinnerer et al. 2023), and the F2100W filter (Hassani et al. 2023). In a future paper we will examine the impact of using different selection strategies in more detail.

Table 1
Census of Massive Young Star Clusters in NGC 1365

ID	CI ID ^a	Other ID	R.A. (J2000) (deg)	Decl. (J2000) (deg)	Log Age ^b (H, H-JW) (Myr)	Log Mass ^c (H, JW) (M_{\odot})	A_V ^d (H, JW) (mag)	CI ₂₀₀ ^e (mag)	Class ^f	Detected ^g <i>B, V, I</i>	Detected ^h PAHs
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)	(10)	(11)	(12)
1	53.400281	-36.142577	..., 6.5	..., 6.2	..., 12.5	2.05	3	..., ..., <i>I</i>	3.3, 7.7, 11.3
2	996	...	53.399891	-36.142575	6.0, 6.0	7.1, 6.6	4.6, 4.4	1.83	2	<i>B, V, I</i>	3.3, 7.7, 11.3
3	1051	J ⁱ	53.400691	-36.142127	6.0, 6.0	7.2, 6.9	4.6, 5.4	1.70	2	<i>B, V, I</i>	3.3, 7.7, 11.3
4	1064	...	53.399386	-36.142085	6.6, 6.6	6.1, 6.2	3.1, 4.3	1.69	2	<i>B, V, I</i>	3.3, 7.7, 11.3
5	53.400338	-36.142072	..., 6.5	..., 6.00	..., 4.9	1.64	2	<i>B, V, I</i>	3.3, 7.7, 11.3
6	53.399056	-36.141991	..., 6.5	..., 5.8	..., ...	1.53	3	..., ..., ...	3.3, 7.7, 11.3
7	1116	M2 ^j , SSC3 ^k	53.399870	-36.141833	6.5, 6.5	6.9, 6.6	1.7, 1.0	1.45	1	<i>B, V, I</i>	3.3, 7.7, 11.3
8	1141	...	53.401402	-36.141753	6.7, 6.7	6.6, 6.5	5.4, 4.6	2.05	2	<i>B, V, I</i>	3.3, 7.7, 11.3
9	1251	...	53.400184	-36.141315	6.0, 6.0	6.9, 6.6	4.6, 6.3	1.88	2	<i>B, V, I</i>	3.3, 7.7, 11.3
10	1259	...	53.402834	-36.141274	6.7, 6.7	6.0, 6.2	4.9, 5.8	1.59	2	<i>B, V, I</i>	3.3, ..., ...
11	53.399771	-36.141127	..., 6.5	..., 6.5	..., 4.5	1.66	2	<i>B, V, I</i>	3.3, 7.7, 11.3
12	1296	M3 ^j , SSC6 ^k	53.399552	-36.141121	6.5, 6.5	6.6, 6.6	1.2, 2.0	1.62	1	<i>B, V, I</i>	3.3, 7.7, 11.3
13	53.402935	-36.140825	..., 6.5	..., 5.9	..., ...	1.39	3	..., ..., ...	3.3, 7.7, ...
14	53.402318	-36.140610	..., 6.5	..., 5.8	..., 3.8	2.69	1	<i>B, V, I</i>	..., 7.7, 11.3
15	53.403212	-36.140610	..., 6.5	..., 6.3	..., 12.2	1.52	3	..., ..., <i>I</i>	3.3, 7.7, ...
16	53.402911	-36.140455	..., 6.5	..., 5.8	..., 8.0	1.73	2	<i>B, V, I</i>	3.3, 7.7, 11.3
17	53.402548	-36.140381	..., 6.5	..., 6.0	..., ...	1.79	3	..., ..., ...	3.3, 7.7, 11.3
18	1506	...	53.403394	-36.140211	6.0, 6.0	6.4, 6.2	4.6, 6.5	1.76	2	<i>B, V, I</i>	3.3, ..., ...
19	53.399524	-36.140090	..., 6.5	..., 6.5	..., 8.4	1.89	3	..., <i>V, I</i>	3.3, 7.7, 11.3
20	1528	...	53.403379	-36.140083	6.5, 6.5	6.2, 6.1	4.6, 6.1	1.74	2	<i>B, V, I</i>	3.3, 7.7, ...
21	53.403331	-36.139860	..., 6.5	..., 6.2	..., 6.9	1.52	2	<i>B, V, I</i>	3.3, 7.7, 11.3
22	53.400461	-36.139849	..., 6.5	..., 6.7	..., 7.4	1.72	2	<i>B, V, I</i>	3.3, 7.7, 11.3
23	53.403374	-36.139568	..., 6.5	..., 6.3	..., 5.6	1.74	2	<i>B, V, I</i>	3.3, 7.7, ...
24	1640	...	53.399837	-36.139432	6.0, 6.0	6.3, 5.8	4.6, 2.4	2.03	1	<i>B, V, I</i>	3.3, ..., ...
25	53.400928	-36.139100	..., 6.5	..., 6.7	..., ...	1.84	3	..., ..., <i>I</i>	3.3, ..., 11.3
26	1700	CI ⁱ	53.404026	-36.139017	6.7, 6.7	6.9, 6.7	4.6, 2.9	1.75	1	<i>B, V, I</i>	3.3, 7.7, 11.3
27	53.402918	-36.138860	..., 6.5	..., 6.6	..., 9.5	1.79	3	..., ..., <i>I</i>	3.3, ..., ...
28	...	M6 ^j , G ⁱ	53.403181	-36.138475	..., 6.5	..., 6.6	..., 19.3	1.85	3	..., ..., <i>I</i>	3.3, SAT, 11.3
29	...	M4 ^j , D ⁱ	53.401702	-36.138433	..., 6.5	..., 6.0	..., ...	1.62	3	..., ..., ...	3.3, SAT, 11.3
30	1782	...	53.400982	-36.138416	6.5, 6.5	6.2, 6.3	4.9, 7.5	1.55	2	<i>B, V, I</i>	3.3, 7.7, 11.3
31	1794	...	53.401669	-36.138367	6.6, 6.6	6.1, 6.0	4.6, 5.8	1.68	2	<i>B, V, I</i>	3.3, 7.7, 11.3
32	1810	...	53.401369	-36.138271	6.9, 6.9	6.2, 6.8	4.6, 7.7	1.97	2	..., <i>V, I</i>	3.3, ..., ...
33	1920	M5 ^j , E ⁱ	53.402514	-36.137680	..., 6.5	..., 6.6	..., 4.3	1.55	2	<i>B, V, I</i>	3.3, SAT, 11.3
34	1932	...	53.402534	-36.137596	6.7, 6.7	6.4, 5.8	5.0, 7.4	1.45	2	<i>B, V, I</i>	3.3, 7.7, 11.3
35	1948	...	53.403066	-36.137494	6.5, 6.5	6.5, 5.8	6.9, 6.6	1.57	2	<i>B, V, I</i>	3.3, 7.7, 11.3
36	53.403115	-36.137151	..., 6.5	..., 6.3	..., 6.4	1.42	2	<i>B, V, I</i>	3.3, 7.7, 11.3
37	2242	...	53.407333	-36.135815	6.6, 6.5	6.2, 6.3	2.0, 2.6	1.71	1	<i>B, V, I</i>	3.3, 7.7, 11.3

Notes.

^a Cluster ID from the NGC 1365 compact cluster catalog, as found at <https://archive.stsci.edu/hlsp/phangs-hst>.

^b Log age—First value: Hubble value from Whitmore et al. (2023). Typical errors are 0.4 dex. Second value: adopted value from first value if available (Hubble), or 6.5 otherwise (JWST—typical errors are 0.6 dex). See Section 3.2.4 for discussion.

^c Log mass—First value: Hubble value from Whitmore et al. (2023). Second value: value from mean of values based on F200W formula (i.e., $\log(M/M_{\odot}) = -0.4068 \times m_{F200W} + 13.83$) and F300M formula (i.e., $\log(M/M_{\odot}) = -0.3639 \times m_{F300M} + 13.25$). Typical errors are 0.35 dex. See Section 3.2.3.

^d Extinction (A_V)—First value: Hubble value from Whitmore et al. (2023). Second value: value from formula $A_V(\text{JWST}) = [(m_I - m_{F360M}) \times 0.810 + 0.99] \times 3.1$. Typical errors are 1.0 mag. See Section 3.2.2.

^e Concentration Index (CI) using the difference in aperture magnitudes in the F200W filter using radii of 1 and 4 pixels. See Section 3.2.1.

^f Visual estimate of how embedded the cluster is. Class 1 = largely unobscured ($A_V \lesssim 3$ mag); Class 2 = partly embedded ($3 \lesssim A_V \lesssim 10$ mag); Class 3 = embedded ($A_V \gtrsim 10$ mag).

^g Brighter than 26th mag in relevant band.

^h Brighter than 20th mag in relevant band.

ⁱ Radio source ID (Sandqvist et al. 1995).

^j MID IR ID (Galliano et al. 2008).

^k SSC = super star cluster ID (Kristen et al. 1997).

Our final census of 37 massive (mass $\gtrsim 10^6 M_{\odot}$), young (age $\lesssim 10$ Myr) candidate clusters in NGC 1365 is compiled in Table 1 and shown in Figure 2. Most of the clusters are in an inner star-forming ring, as discussed by Schinnerer et al. (2023), while a few more are in the bar just to the NE of this ring (i.e., objects 33 through 37).

We find that NGC 1365, at a distance of about 20 Mpc, has the richest population of massive young clusters of any known galaxy within 30 Mpc, with about 30 clusters in this mass and age range compared with roughly a dozen in the merging Antennae galaxies (Whitmore et al. 2010), located at ~ 22 Mpc. More distant galaxies with comparable populations are

NGC 3256 (65 clusters, distance of 38 Mpc) and NGC 3690 (27 clusters, distance of 42 Mpc), according to Adamo et al. (2020a). Both of these are ULIRGs and ongoing merging systems, with estimated star formation rate = $\text{SFR} \approx 40 M_{\odot} \text{ yr}^{-1}$, roughly twice as large as in NGC 1365 (Lee et al. 2022).

We are now ready to address the question: “how many massive young clusters are so enshrouded in dust that they are missing at optical wavelengths?” From column (11) of Table 1 we find only 4 out of 37 (11%) of our cluster catalog are so dusty that they are not clearly visible in the *I*-band image (longest wavelength observed from HST). However, age and mass estimates are quite challenging without photometry in at least three optical bands, which means clusters that are not also detected in the *B* and *V* filters are unlikely to be included in optically selected cluster catalogs, and hence might be considered missing in some cases. For NGC 1365, we find that 9 of 37 ($\sim 24\%$) clusters are “missing” in the *V* band and 11 of 37 ($\sim 30\%$) are “missing” in the *B* band. We therefore find that between $\sim 11\%$ – 30% of the clusters are “missing” from optical catalogs because they are embedded at some level. This is similar to the situation in the Antennae, where only $2/13 = 15\%$ of bright thermal radio continuum sources were “missing” at optical wavelengths, and 38% of fainter clusters (with $S > 70 \mu\text{Jy}$) were missing (Whitmore & Zhang 2002).

Figure 3 reveals a number of interesting features in Region 1. The top left panel shows the strong dust lane cutting through an optical *IVB* image taken with HST. Three different JWST +HST filter combinations (as indicated along the upper right corner of each panel) reveal embedded young clusters, some of which are very bright and saturated (e.g., M4, M5, and M6 in the F2100W and F770W images), as well as some blue, slightly older clusters outside the dust lane. All three of these saturated clusters are also strong radio continuum sources (corresponding to sources D, E, and G, respectively, in Sandqvist et al. 1995).

There are also very thin, dark filaments visible in the F335M/F200W/*I*-band image; a few are labeled in Figure 3. These are typically found along the inside edge of some of the strongest F335M PAH ridges, and have widths of a few tens of parsecs and lengths of a few hundred parsecs, similar to the massive filaments found in the Milky Way (Hacar et al. 2022), normal spiral galaxies (e.g., Thilker et al. 2023; Meidt et al. 2023), as well as other star bursting galaxies such as the Antennae (Whitmore et al. 2014).

3.2. Estimate of Star Cluster Properties

Figure 4 presents four diagnostic diagrams that will be used in this section to estimate several basic properties of the clusters. These include a color–magnitude diagram, a color–color diagram, and two flux versus concentration index diagrams using different bands. A visual estimate of the level of embeddedness is included in Table 1: Class 1 = largely unobscured ($A_V \lesssim 3$), Class 2 = partly embedded ($3 \lesssim A_V \lesssim 10$), and Class 3 = embedded ($A_V \gtrsim 10$). The location of data points for three clusters is shown in each panel to help demonstrate the properties of a largely unobscured cluster (M3, shown in blue), a partially embedded cluster (obj 34, shown in orange), and an embedded cluster (M6, shown in red). Arrows show how properties track in the diagram, including PAH emission strength (youngest clusters), extinction (dustiest clusters), and size (to help separate stars from barely resolved clusters).

We create several different training sets to illustrate a number of features. These include the 37 massive young clusters discussed in Section 3.1, now broken into three subsamples: unobscured, partly embedded, and embedded. Other training sets include older clusters (with ages around 300 Myr), a set of largely embedded F200W sources in the dust lane around M4 (see Figure 3), and objects that are likely to be red super giants (RSGs) in a stellar association near the eastern end of the bar. The color coding for the different samples is shown in the top left panel of Figure 4.

In the next subsections, we show how we separate individual stars from clusters and then describe methods to estimate the extinction (Section 3.2.2), mass (Section 3.2.3), and age (Section 3.2.4) of embedded clusters. We adopt the estimated properties of relatively unobscured clusters determined in Whitmore et al. (2023) using Hubble data, as included in Table 1, to establish correlations that allow us to estimate properties for the more deeply embedded clusters observed primarily in the JWST bands.

3.2.1. Separation of Clusters and Stars

The bottom left and bottom middle panels in Figure 4 show that the F200W filter provides a powerful new way to differentiate stars from barely resolved clusters and to eventually estimate the sizes of the clusters. The scatter in the concentration index for stars (black points) using the F200W filter on JWST is dramatically less than from the *I* filter on HST. If we use the *V* filter from HST instead the results are intermediate between F200W and *I*, but we are not able to penetrate the dust as well. This better resolving power might be expected since the resolution with JWST in the F200W band is approximately $0''.05$, while it is approximately $0''.08$ in *I* using Hubble. In addition, the sampling is better with $0''.031 \text{ pix}^{-1}$ for the F200W filter compared to $0''.040 \text{ pix}^{-1}$ for the *I* filter, and the former provides these measurements even in dusty regions.

Figure 4 shows that while there is a large overlap between stars and clusters in CI_I , there is no overlap using the CI_{F200W} measurements, and the orange line is able to cleanly separate stars from clusters in all cases for our training sets. In the future, we plan to use this new capability to improve the size measurements of star clusters as well.

This improved performance also shows that while some of the embedded F200W objects (green points) in the bottom left panel of Figure 3 are likely to be stars, most (particularly the brighter ones) appear to be clusters, probably with mass in the range 10^5 – $10^6 M_{\odot}$.

3.2.2. Extinction Estimates for Embedded Clusters

There are two approaches we can take to estimate extinction. The first is to use the Galactic extinction law from (Fitzpatrick 1999). The *I*–F360M versus F300M–F335M diagram shown in the upper right panel of Figure 4 demonstrates the range of extinction (A_V) values found in our sample. A reddening vector following a Galactic extinction law is shown for $A_V = 10$ mag. We find that this predicts 4.75 mag of reddening in *I*–F360M, hence the coefficient connecting A_V and *I*–F360M is $10/4.75 = 2.11$. Very small deviations from exactly vertical are ignored since the horizontal dimension is primarily an indication of PAH emission strength rather than reddening. For example, both M6 and obj 34 are found in

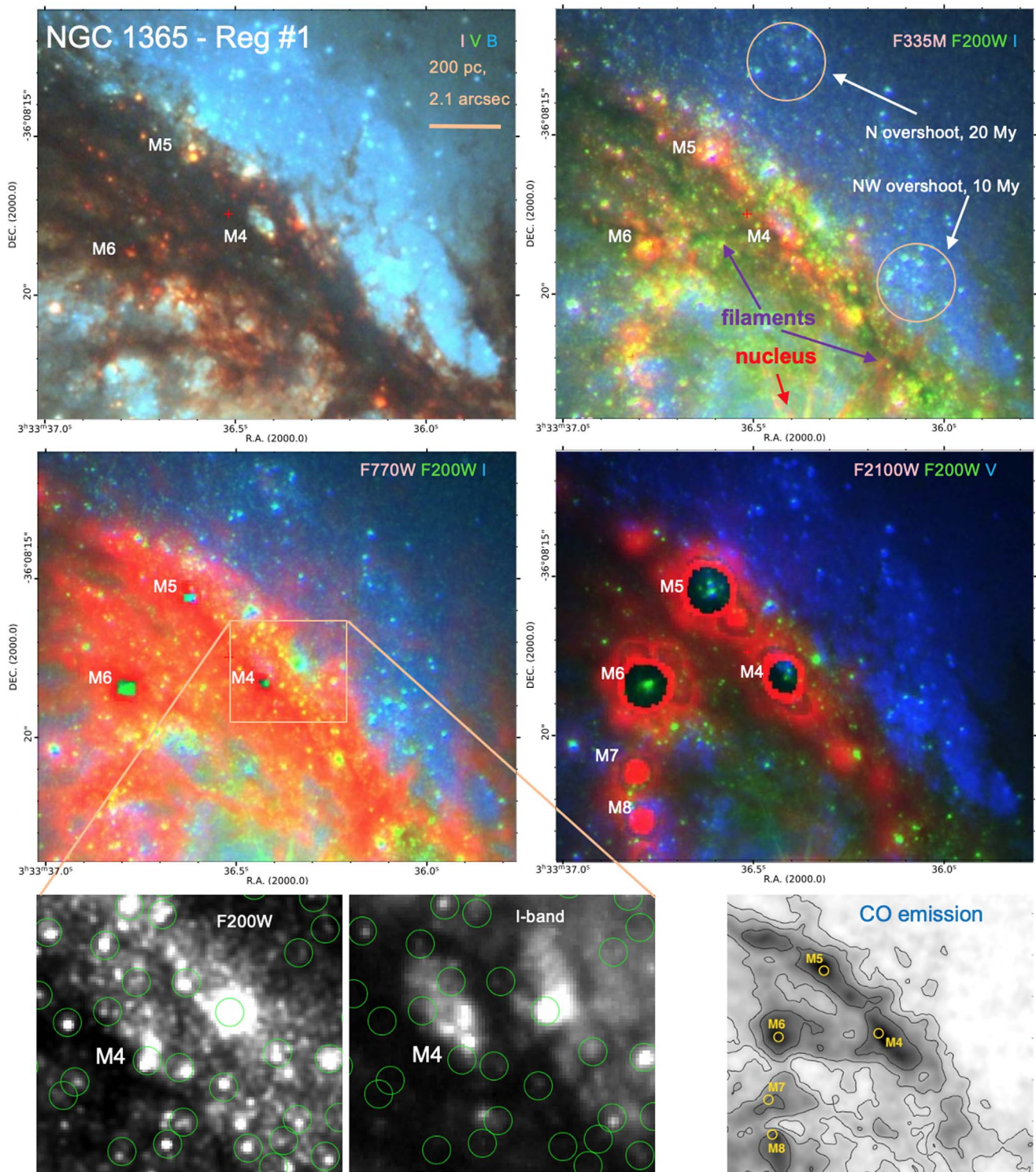


Figure 3. Region 1 in NGC 1365. The figure shows an *IVB* image from Hubble in the upper left, and three images with different JWST bands (F335M, F700W, or F2100W) in red; F200W (NIR continuum) in green; and *I* or *V* from Hubble in blue in the other panels. This region is just north of the nucleus and contains the three strongest radio continuum sources in this galaxy, which appear to be young star clusters with masses of roughly \log mass in the range 6–7 (referred to as M4, M5, and M6 by Sandqvist et al. 1995 and Galliano et al. 2008, 2012). Several thin dark filaments are shown along a ridge line containing some of the most recent star cluster formation in the F335M image. The holes in M4, M5, and M6 in both the F2100W and F770W images show where the data are saturated. The largely unobscured region to the upper right shows an “overshoot region” outside the inner part of the galaxy (see Sormani et al. 2020), with most of the clusters having estimated ages in the range between 10 and 20 Myr. The two bottom left panels show the excellent resolution of the JWST F200W image and its ability to cut through the dust better than the *I*-band Hubble image, which is shown to the right. The bottom right panel shows the CO (2–1) line intensity emission map (Schinnerer et al. 2023) for Region 1. The small red cross in the top four panels is an artifact of the Hubble Legacy Archive, which we used to make the figures, and should be ignored. Linear diagonal artifacts from the saturated nucleus are visible at the bottom edge of the three color panels using JWST filters.

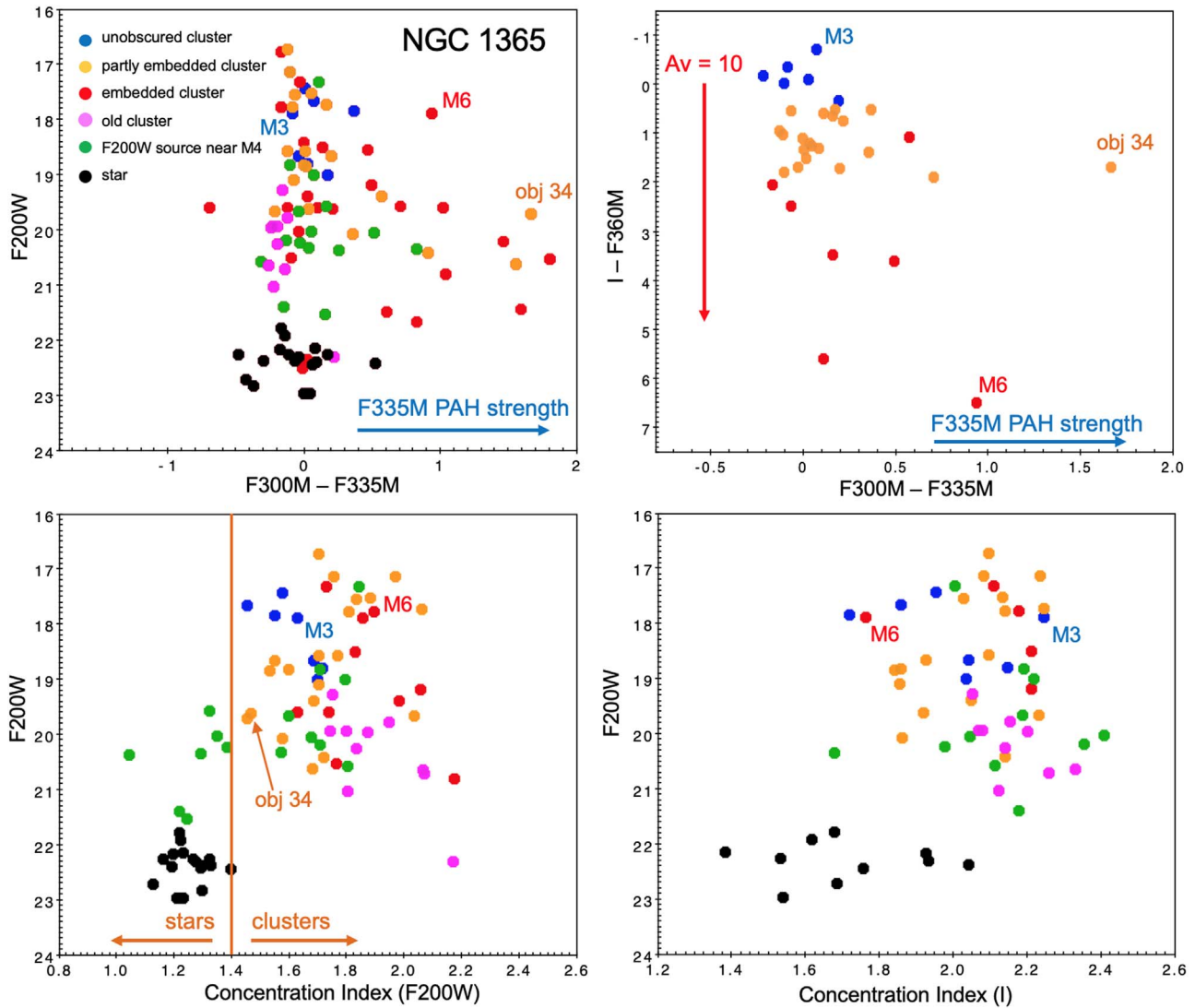


Figure 4. Diagnostic diagrams for clusters and stars in NGC 1365. Top left: F300M–F335M vs. F200W color–magnitude diagram including unobscured clusters (blue circles), partly embedded clusters (orange), embedded clusters (red), old clusters (pink), F200W sources near M4 (both clusters and stars; green), and stars (black) from the outer part of the galaxy. Top right: F300M–F335M versus I –F360M color–color diagram, using the same color coding. Bottom left: concentration index measured from the F200W filter. The vertical orange line shows that stars and clusters nicely separate at $CI_{F200W} = 1.4$. Bottom right: the same diagram as to the left, but now using CI measurements in the I filter, which do not distinguish stars from clusters as well as the JWST F200W filter.

regions with strong F335M emission, explaining their positions on the right side of both the color–magnitude and color–color diagrams. We find that the majority of the clusters have A_V values less than 10 mag. The underlying assumption needed to estimate extinction using the I –F360M color index is that all the (young) clusters in our sample have similar intrinsic colors; hence, the redder color is caused by extinction rather than due to age. The fact that unobscured clusters (blue points) all have I –F360M values around 0 supports this assumption.

A second approach is to use the empirical correlation between the $E(B - V)$ values from the spectral energy distribution (SED) -fitting based on Hubble data versus I –F360M, as shown in the bottom panel of Figure 5. We use the sample of Class 1 and 2 (symmetric and asymmetric clusters, human classified sample; see Whitmore et al. 2021) in NGC 1365 to make this determination. Turner et al. (2021) describe the basic SED age-dating method used in the PHANGS–HST pipeline. More recently, Whitmore et al. (2023) described modifications that identify and correct the age estimates for old

globular clusters in the pipeline; many old globular clusters have incorrect age estimates in the default pipeline due to challenges in disentangling age/reddening/metallicity. This helps remove interlopers that might otherwise pollute our young clusters sample. The fraction of clusters with incorrect age estimates in NGC 1365 are reduced from $\approx 20\%$ to $\approx 10\%$ according to Whitmore et al. (2023).

Only the log age < 6.7 clusters are included to make the sample similar to our program sample of young massive clusters in NGC 1365. The rms scatter (0.95 mag) is relatively large and the correlation coefficient (0.49) is only fair; hence, our estimates of extinction will only be approximate. The resulting formula is $A_V(\text{JWST}) = [(m_I - m_{F360M}) \times 0.810 + 0.99] \times 3.1$, which provides estimates of extinction for our partially embedded clusters. Hence, the coefficient connecting A_V and I –F360M is $0.810 \times 3.1 = 2.51$, relatively similar to the estimate from the Galactic extinction law discussed above (i.e., 2.11). This is the formula we will use in Table 1.

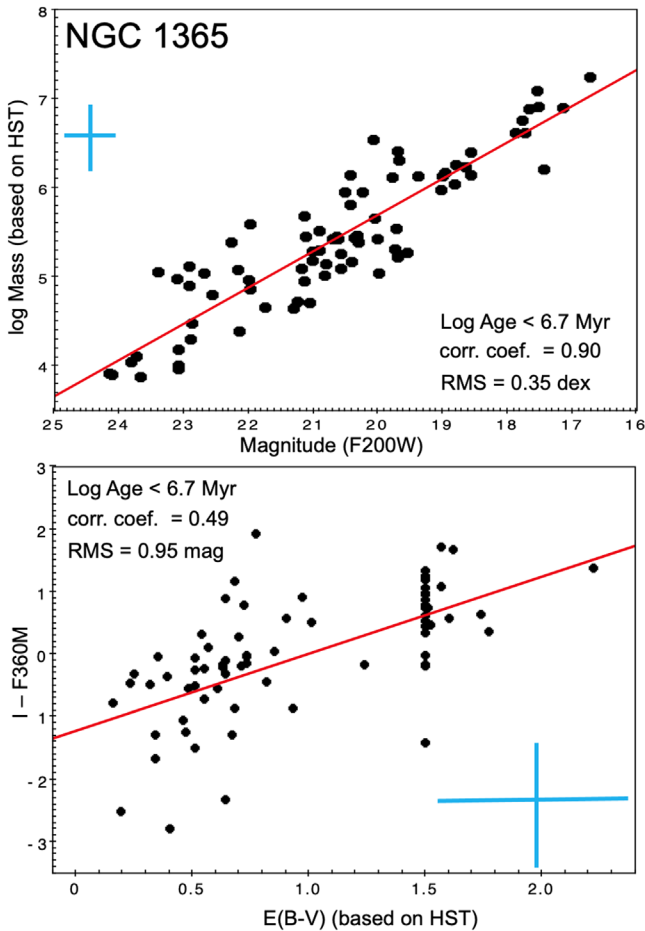


Figure 5. The top panel shows the correlation between the JWST F200W magnitudes and the HST-based cluster mass estimates from Whitmore et al. (2023) for Class 1 (symmetric) and 2 (asymmetric) clusters for a sample with log age < 6.7 . The bottom panel shows the relationship between $I - F360M$ colors and the HST-based reddening ($E(B - V)$) estimates for the same sample. The correlation coefficient and the rms scatter are provided, and the crosses show the typical uncertainty.

In principle it would be advantageous to use only JWST observations to determine the correlation, since four of our clusters do not have I -band photometry. However, we found that the smaller baseline resulted in considerably more scatter, and hence was not practical. We note that the relationship we derive is specific to the analysis presented here since aperture corrections have not been made to our data yet; other researchers will need to determine the constants from their own photometry.

Table 1 shows that the highest estimated extinction value is for the intrinsically bright object 28 (=M6/G, e.g., saturated in the F770W and F2100W filter observations, as shown in Figure 3), which has an estimate of $A_V = 19.3$ mag. The value would be about $A_V = 15$ mag if the Galactic extinction formula was used instead. This extinction is even larger than WS80, a deeply embedded star cluster in the Antennae galaxies (Whitmore & Zhang 2002) that is the strongest CO and $15 \mu\text{m}$ source in that system, and has an estimated age of 2 Myr, $A_V \approx 8$ mag, and log mass ≈ 6.6 . This indicates that the very different environments of a chaotic merger and the central region of a barred spiral galaxy can both produce very massive, heavily extinguished, young clusters.

3.2.3. Mass Estimates for Embedded Clusters

To estimate masses for our embedded clusters we use the approach developed by Rodriguez et al. (2023), who show that there is a good correlation between the F200W flux and the HST-based mass estimates for clusters in NGC 7496, when limited to very young clusters. A similarly good correlation is seen for F335M. This is because similar age clusters have similar mass-to-light ratios.

As shown in the top panel in Figure 5 we find a very good correlation in NGC 1365 as well, with a correlation coefficient of 0.90 and an rms scatter of just 0.35 dex. The resulting formula is $\log(M/M_\odot) = -0.4068 \times m_{F200W} + 13.83$. We again use the sample with log age < 6.7 clusters to make it similar to our sample of young massive clusters in NGC 1365. A version of this fit based on the F300M magnitude is nearly as good: $\log(M/M_\odot) = -0.3639 \times m_{F300M} + 13.25$, with correlation coefficient = 0.87. If the mass-to-light ratios were exactly equal for all the objects we would find the coefficient in the formulae would be the inverse of 2.5 in the definition of magnitudes (i.e., $-2.5 \log \text{flux}$); hence, it would be $1/2.5 = 0.40$. We note that this is true with our empirically determined coefficients approximately equal to 0.41 and 0.36. The final mass values included in Table 1 use the mean of these two mass estimates.

Our estimates for clusters M4, M5, and M6 range from log mass = 6.0 to 6.6, in fair agreement with Galliano et al. (2008), who estimate values “on the order of $10^7 M_\odot$ ” for the three clusters. We note that in Galliano et al. (2012), the estimates are slightly lower, ranging from log mass = 6.5 to 7.0, which are in better agreement with our values.

We find that 7 of the 37 clusters in Table 1 have estimated masses that are somewhat below log mass = 6.0. This is because there are uncertainties at the factor of $\sim 2-3$ level associated with each cluster mass. Going forward, for various estimates we assume that 30 clusters from our sample have masses of at least log mass $\gtrsim 6.0$.

3.2.4. Age Estimates for Embedded Clusters

In the future, we will combine HST and JWST bands to improve age estimates, especially for partially embedded clusters. For now, we will use the same basic approach as in the previous two subsections to estimate the ages of the embedded and partly embedded clusters based on HST estimates.

As shown in Figure 6, the age estimates are somewhat more complicated than extinction and mass estimates, with more scatter and a nonlinear dependence. This figure plots HST-based ages versus source brightness in the F335M filter, which includes the $3.3 \mu\text{m}$ PAH feature and also continuum emission. While there appears to be a possible correlation for ages between log age = 6 and 8.3 (e.g., correlation coefficient = 0.44), there is essentially no correlation at older ages.

However, at the bright end (i.e., with F335M < 20.0 mag) we find that most sources are young, with log age less than 7. A visual examination of the clusters in the upper right part of the diagram show that five of the six (circled) are in regions with strong $H\alpha$ emission and are likely to be young; hence, the HST-based age estimates are likely to be incorrect. An example is M5 (object 28), which is a very strong radio source that has been age dated using spectral observations indicating that it is

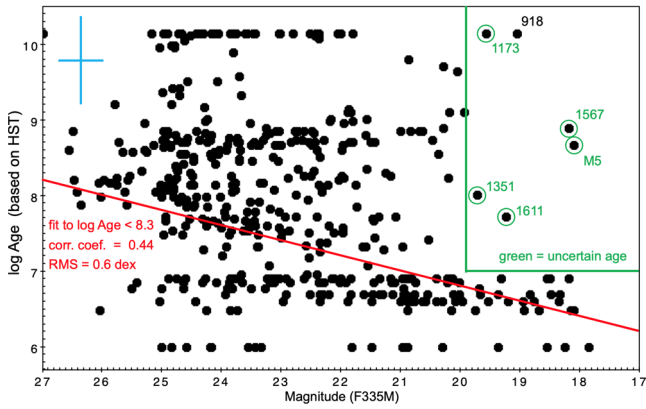


Figure 6. A plot of JWST F335M magnitude vs. age estimates based on Hubble data from Whitmore et al. (2023) for Class 1 (symmetric) and 2 (asymmetric) clusters. This shows a large degree of scatter and essentially no correlation for clusters with F335M brighter than 20 mag have log age values less than ≈ 7 (10 Myr). A visual examination of the objects in the upper right rectangle shows that all but one of the clusters (i.e., the objects with green circles around them) have $H\alpha$ emission nearby, indicating that the older age estimate is probably incorrect. The blue cross shows the approximate uncertainty based on a fit to the points with log age ≤ 8.3 , shown as the red line.

clearly younger than 10 Myr (Galliano et al. 2012). We note that the plots for other PAH bands, notably F770W and F1130W, look very similar to that for F335M.

Based on this evidence, we believe at least 90% of the objects with F335M brighter than 20 mag are younger than 10 Myr. We note that this would be compatible with the results from Whitmore et al. (2023), who find that approximately 10% of the HST-based age estimates of clusters in the inner region of NGC 1365 are overestimated, due to the large amount of dust that makes the clusters appear red and potentially old. While low-resolution ($0''.8$) $H\alpha$ imaging from MUSE (Emsellem et al. 2022) has been used in Whitmore et al. (2023) to improve the age-dating to some extent, this is not as good as having high-resolution Hubble $H\alpha$ imaging.

Figure 7 shows there are physical reasons why we expect most of the brightest PAH emitters to be young. This figure shows a normalized (at 1 Myr) plot of predicted flux values in various MIRI bands from the Draine models of interstellar dust emission (Draine et al. 2021). There are similar profiles for the NUV and FUV emission that are the main driver of the PAH emission. In all these cases the peak flux occurs for ages around 2 to 4 Myr, with relatively rapid declines within just 10 Myr.

As expected, the nebular emission (purple line) shows an even more rapid drop-off, falling to just a few percent by 6 or 7 Myr. This is why nebular lines, such as $H\alpha$, are so valuable for age-dating; if you see strong emission associated with a cluster, the source is young (e.g., Whitmore & Zhang 2002; Anders & Fritze-v. Alvensleben 2003; Chandar et al. 2010; Foesneau et al. 2012). The PAH band strengths have a similar predicted time evolution, although they typically drop to only $\sim 20\%$ of their initial flux by 10 Myr. This is why the PAH emission is correlated with young cluster ages in Figure 6. We note that this drop-off in PAH strength as a function of time also appears to be compatible with the results of Lin et al. (2020) for LEGUS galaxies.

Since the scatter in the relationship between F335M magnitudes and log age based on HST is fairly large (~ 0.6 dex), we have chosen to simply assign log age values of 6.5 (3 Myr) for the clusters without HST age estimates. This is

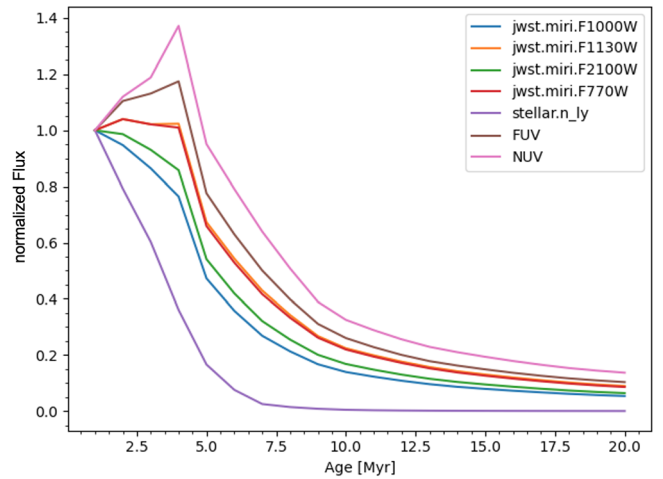


Figure 7. The predicted evolution of the flux emitted from clusters in a variety of filters, from CIGALE (Boquien et al. 2019). We include the JWST MIRI F770W, F1000W, F1130W, and F2100W filters, each including contributions from the nebular emission, which is an important component (i.e., without the nebular emission the various MIRI bands are nearly on top of each other); the NUV and FUV filters from GALAX; and the nebular emission alone (i.e., proportional to $Ly\alpha$ in this case but similar to other line emission such as $H\alpha$). A legend is provided in the upper right. The profiles are normalized to 1.0 at an age of 1 Myr. The key point is that strong emission in the JWST bands is generally indicative of young ages, but the drop-off is not as steep a function of age as for the nebular lines.

consistent with the results from Figure 6, the flux profiles in Figure 7, and the comparison with Hubble age estimates in Table 1. Hence, our final age estimates are included as the second entry in column (6) in Table 1 and are either the HST-based SED-fitting value if it exists or log age = 6.5 (3 Myr) for the clusters without estimates from Hubble.

3.3. Duration of the Embedded Phase

A complete census of massive clusters from the optical and infrared allows us to estimate the time that clusters remain in the deeply embedded phase, specifically the length of time they are not detectable at optical wavelengths. While early estimates (e.g., Kawamura et al. 2009) were longer than 10 Myr, more recent values have generally been in the range 1–5 Myr (e.g., Lada & Lada 2003; Whitmore et al. 2014; Hollyhead et al. 2015; Grasha et al. 2018; Chevance et al. 2020; Kim et al. 2021; Hannon et al. 2022; He et al. 2022; Kim et al. 2023). NGC 1365 provides a particularly interesting case because of the large amount of gas and dust that might, in principle, result in longer timescales. It also provides a unique view of the embedded phase of extremely massive ($\geq 10^6 M_\odot$) clusters, which can eventually be compared with results from lower-mass clusters to assess whether or not the duration of this phase has any dependence on cluster mass.

We calculate the duration of the deeply embedded phase for massive clusters in NGC 1365 as No. missing / total \times age (i.e., sample interval) = $4/30 \times 10$ Myr = 1.3 ± 0.7 Myr, a very short timescale. If we instead consider the timescale for massive clusters to be detectable in the B , V , and I bands (a minimum needed for most optically based age-dating methods), we find a somewhat longer timescale No. missing / total \times age = $11/30 \times 10$ Myr = 3.7 ± 1.1 Myr.

Our estimates for how long very massive clusters in NGC 1365 remain embedded in their natal interstellar medium (ISM) clouds are consistent with those found by previous

works, without being obviously longer or shorter relative to estimates made for lower-mass clusters in other galaxies. It is important to establish if the length of the embedded phase depends on cluster mass or not, because this provides direct constraints on early feedback processes that may be important for limiting the masses of clusters.

4. Region 2: Star Formation Triggered by Collisions of Gas and Dust Streamers

4.1. Triggered Star Formation

While the strongest star formation in NGC 1365 is occurring in the inner region (i.e., 36 of the 37 youngest most massive clusters as we saw in Section 3.1), a moderate amount is also taking place along the bar, an example of which is shown in Figure 8. In particular, Bubble 1 is object 37 from Table 1, with a $\log \text{mass} = 6.6$. Much of the star formation in Region 2 appears to be triggered by dust and molecular gas from dark filamentary “streamers” that reach the bar, colliding and merging with the gas and dust as illustrated in the bottom panel of Figure 8, and as discussed in Elmegreen et al. (2009). Another excellent candidate for this type of triggered cluster formation is Bubble 2 (3 Myr, $A_V = 2.1$ mag; age and extinction estimates from Whitmore et al. 2023), as shown in Figure 8. A careful look at Figure 1 shows several other examples of regions of star formation near the intersection of streamers and the bar.

Previous kinematic work has shown strong streaming motions due to the bar in NGC 1365 (e.g., Sakamoto et al. 2007; Lindblad et al. 1996; Schinnerer et al. 2023). Simulations provide a more visual way to understand the dynamics of streamers of gas as they approach, interact with, and sometimes overshoot the bar lane. We encourage interested readers to watch this video.²⁸ Figure 11 in Sormani et al. (2020) also provides insight to understand the dynamics and the role of streamers as they interact with the bar, potentially triggering star formation. One key new result from JWST observations is how ubiquitous these streamers are in spiral galaxies (see Meidt et al. 2023; Thilker et al. 2023).

Figure 9 shows a larger field of view around Region 2 (i.e., “Region 2 extended” from Figure 1) and contains additional kinematic information that can be used to help test the triggering hypothesis. A full kinematic analysis is beyond the scope of the current Letter, however. The top panel shows an *IVB* Hubble image that highlights the dust lanes and regions of young stars. Besides the streamer we show in Figure 8, which is labeled Streamer 1, we find several other candidate streamers. The second panel uses the F770W PAH emission filter (shown in the red) to highlight the youngest star-forming regions. The third panel shows CO (2–1) moment-1 velocity from the PHANGS–ALMA observations (Leroy et al. 2021), color-coded from -220 km s^{-1} (dark blue) to $+220 \text{ km s}^{-1}$ (dark red and white). CO clouds are marked using ellipses and box-circle symbols. Ellipses are sized to represent the deconvolved major/minor cloud axes (FWHM), whereas box-circles indicate unresolved clouds (E. Rosolowsky et al. 2023, in preparation).

We first note the large number of CO clouds in Streamer 1, indicating that much of the gas and dust feeding into the bar, and eventually into the inner region of the galaxy, is coming

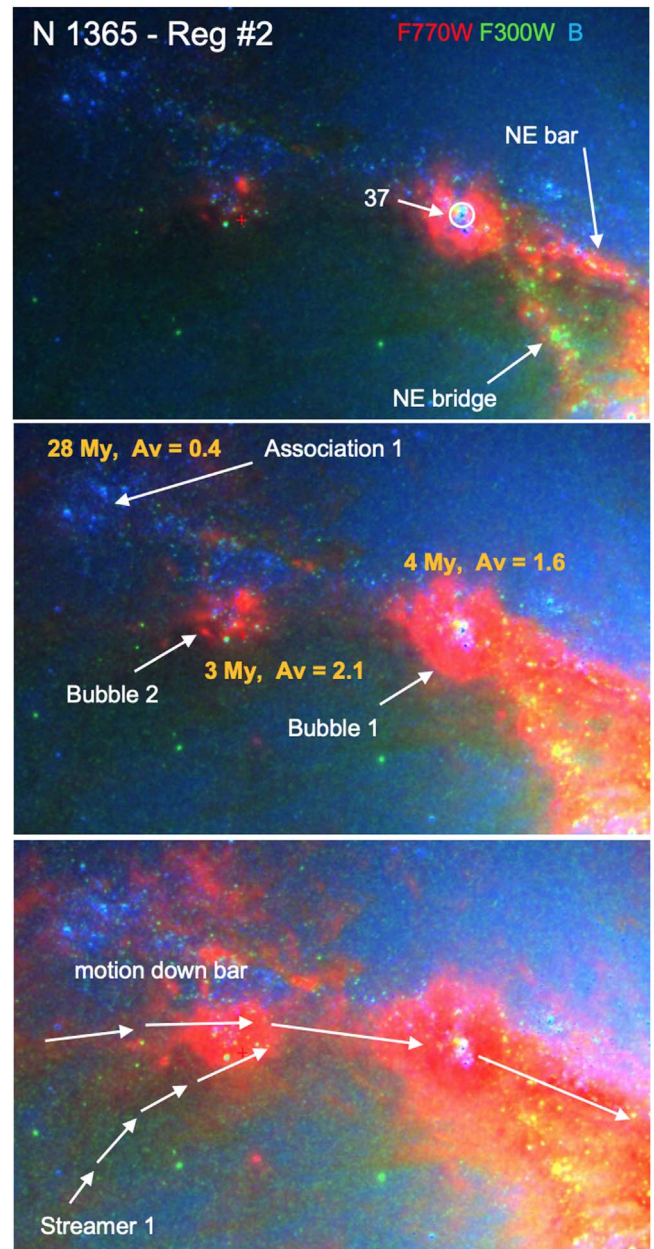


Figure 8. JWST (F770W and F300M) and Hubble (*B*-band) images of Region 2, just up the NE bar from the nucleus, as shown in Figure 1. The different contrast levels highlight key features, including two young “bubbles” resulting from star formation that may have been triggered by an intra-bar streamer colliding with gas and dust in the bar itself, as shown by the arrows. Another feature, which is only visible in the JWST bands due to the heavy dust in the area, is the “NE Bridge,” as discussed in the text.

from the streamers rather than from the bar itself. This figure shows that the gas velocities along Streamer 1 are similar (i.e., this feature appears as a uniform, green color, indicative of velocities around 0 km s^{-1}), as expected for a coherent physical component.

Gas velocities along the bar (the upper blue arrows) are more mixed, with values $\approx +100$ to $+200 \text{ km s}^{-1}$ along most of the bar, but with velocities $\approx +50 \text{ km s}^{-1}$ in regions where the streamers intersect the bar. This appears to be what is happening in the region of Bubble 2, with the velocities shown by a light blue-green color indicating a velocity that is intermediate between that of the bar and Streamer 1, as we

²⁸ <https://www.youtube.com/watch?v=1eJ81GIII&list=PL1sb6ZGKWb16QOFpd6OYU4zLXCmBa6uux&index=5>

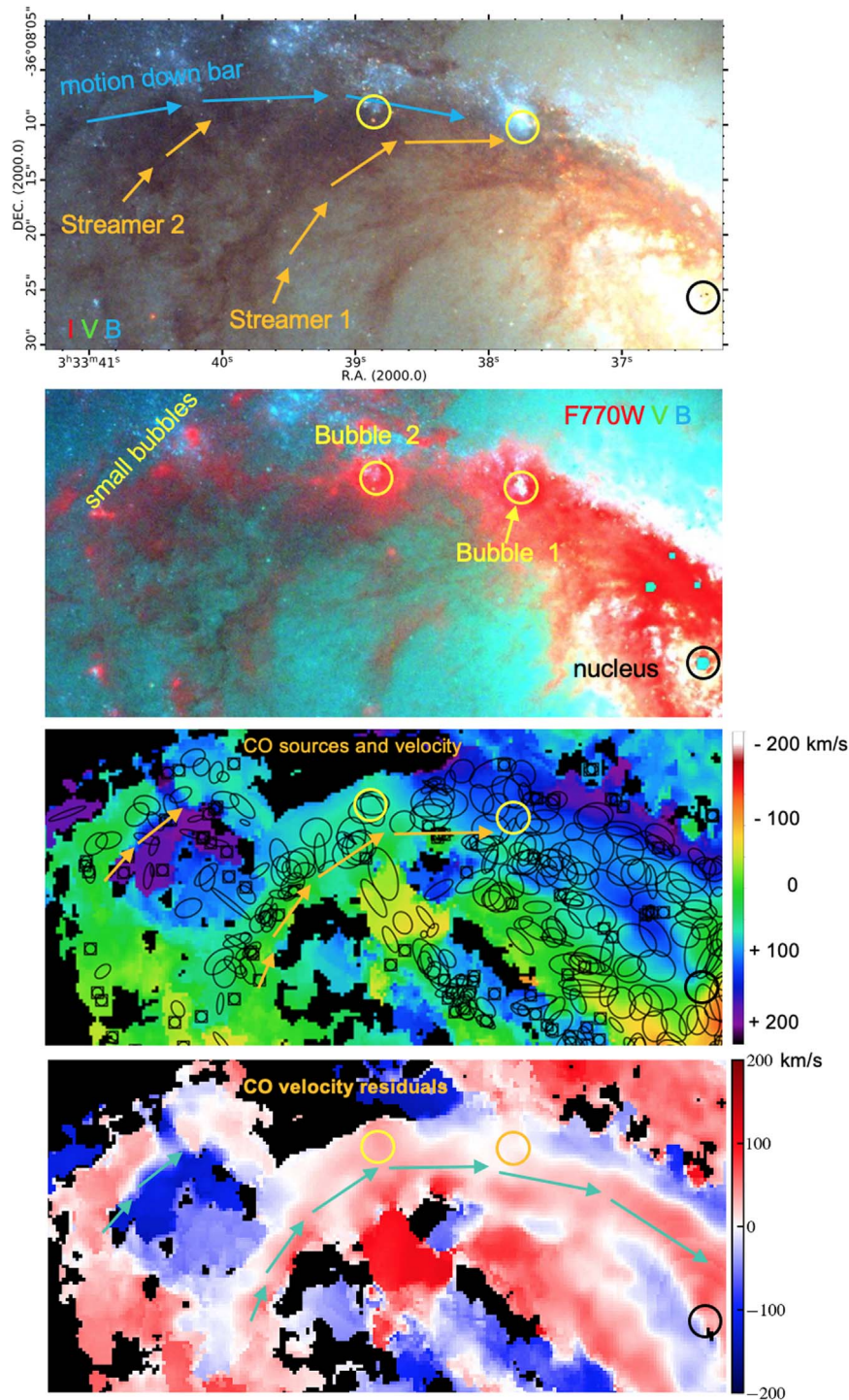


Figure 9. Four images showing the extended Region 2 field of view. The top panel shows *IVB* observations from Hubble, and the second panel shows the combined JWST (F770W) and Hubble (*V* and *B*) images. Objects identified in Figure 8 and the location of Streamer 2 and the nucleus are included. The third panel shows giant molecular clouds identified via their CO(2–1) emission (ellipses represent their fitted spatial extents), while the bottom panel shows velocity residuals from a model of the inner disk rotation from Schinnerer et al. (2023). See the text for discussion.

would expect if this formed from the collision of the two components. The situation is less clear for Bubble 1, but the color appears to be light blue ($\approx +100 \text{ km s}^{-1}$), as expected if it is a combination of dark blue (just to the north and south) and green from Streamer 1.

We also find a less CO-populated streamer farther out in the bar, which we have labeled Streamer 2. It appears to be associated with three or four small star formation regions rather

than a large event like the formation of Bubbles 1 and 2. However, we find that it has a similar morphology with uniform green colors until it mixes with the bar and then shows a combination of green and blue velocities beyond that point.

The bottom panel of Figure 9 shows the CO velocity residuals based on a model of the inner rotating disk (Schinnerer et al. 2023). This supports the same basic interpretation, but provides an even stronger connection

between Streamer 1 and the feeding of the inner region, with a relatively uniform, narrow red component all the way into the central region. We also note that the association of Bubble 1 with the streamer is clearer in the residual velocity image than the original velocity map just above it.

4.2. Clues to Orbital Dynamics in NGC 1365: The NE Bridge and Other Overshoot Regions

An important capability of JWST is the ability to penetrate through the dust and identify structures that are not observable at optical wavelengths. An example is the NE bridge, which is identified in the top panel of Figure 8, but is much less obvious in the HST image in Figure 1. Features like this bridge are often seen in hydrodynamical simulations of bars (e.g., Sormani et al. 2020; Tress et al. 2020; Henshaw et al. 2022; Schinnerer et al. 2023).

Gas in a barred potential travels almost radially along the bar dust lanes (in the frame corotating with the bar) toward the inner parts of the galaxy. However, not all the gas transitioning from the dust lanes to the inner regions does so smoothly—some overshoots and collides with gas traveling down the opposite dust lane.

We believe the NE bridge feature is a clear example of material that has overshoot the inner ring after falling toward the center along the southern bar dust lane (bottom right of Figure 1). When the material reaches the inner ring, gas and stars behave differently because they obey different equations of motion. Gas is affected by cloud collisions and thus dissipates some of its energy and eventually ends up accreting onto the inner ring. Stars, on the other hand, do not feel pressure forces and are unaffected by cloud collisions, so they fly undisturbed through the gas clouds and continue on their ballistic orbits toward the bar lane on the opposite (northern) bar lane, as appears to be the case with the NE bridge.

These, and several other features obvious in the new JWST observations (e.g., the overshoot regions marked in Figures 3 and 2), provide important insights into how the orbits of stars and gas vary depending on where they form and how this results in the population of older clusters outside the inner star formation ring. A related phenomenon is discussed in Whitmore et al. (2023; i.e., the “300 Myr” overshoot region, as identified in Figure 2).

5. Summary and Conclusions

In this work, we have presented new infrared JWST NIRCAM and MIRI imaging of the highly star-forming barred spiral galaxy NGC 1365 in the F200W, F300M, F335M, F360M, F770W, F1000W, F1130W, and F2100W filters. These images show not only the same compact star clusters and individual stars observable in optical Hubble images, but also reveal recently formed clusters that remain deeply embedded in gas and dust, PAH emission, and new (likely dynamical) structures like the NE bridge and dark thin filaments. From measurements of the concentration index we found that the F200W filter on JWST gives more accurate size measurement for stellar sources than the *I* filter on Hubble, resulting in a cleaner separation of stars and clusters.

We have compiled a comprehensive catalog of 37 massive ($M \gtrsim 10^6 M_\odot$), young (age $\lesssim 10$ Myr) clusters in NGC 1365, all but one located in the $28'' \times 26''$ (2.7 kpc \times 2.5 kpc) region around the central star-forming ring. We found that 4 of 30

($\sim 13\%$; i.e., including only the 30 objects with measured values below $10^6 M_\odot$ in Table 1) are so deeply embedded that they cannot be easily seen in the optical images taken by HST, even in the *I* filter, while 11 of 30 ($\sim 37\%$) do not have mass estimates from HST because they are missing in the HST *B*-band images, and hence their ages cannot be easily determined in the visible using SED-fitting techniques. These statistics imply that the massive clusters in NGC 1365 remain deeply embedded (not detectable in the *I* band) for 1.3 ± 0.7 Myr and sufficiently embedded that they are not detectable in the *B* band for 3.7 ± 1.1 Myr. Hence, we find that NGC 1365 contains the richest known population of massive young clusters in any galaxy within 30 Mpc, even larger than the populations formed in the merging Antennae galaxy.

We also identified several interesting new features near the intersection between the bar and streamers of gas and dust that appear to be colliding and triggering star formation. Other new features include an NE bridge, which may be created by material that dynamically overshoot the star-forming ring after falling toward the inner regions along the dust lane in the southern bar.

The age, mass, and extinction estimates in this Letter are approximate and preliminary, since they are based on correlations between optical estimates and JWST observations for the more embedded clusters. In the future we expect to obtain better age estimates using SED-fitting techniques using both the HST and JWST observations.

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Facilities: HST, JWST, ALMA, VLT-MUSE.

Data Availability

The imaging observations underlying this article can be retrieved from the Mikulski Archive for Space Telescopes at https://archive.stsci.edu/hst/search_retrieve.html under proposal GO-15654. High-level science products, including science ready mosaicked imaging, associated with HST GO-15654 are provided at <https://archive.stsci.edu/hlsp/phangs-hst>. The specific PHANGS-JWST observations analyzed can be accessed via doi:10.17909/9bdf-jn24 and PHANGS-HST images accessed via Lee (2021).

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